

# A Cosmological Unicorn Solution to Finsler Gravity

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We present a new family of exact vacuum solutions to Pfeifer and Wohlfarth’s field equation in Finsler gravity, consisting of Finsler metrics that are Landsbergian but not Berwaldian, also known as unicorns due to their rarity. Interestingly we find that these solutions have a physically viable light cone structure, even though in some cases the signature is not Lorentzian but positive definite. We furthermore find a promising analogy between our solutions and classical FLRW cosmology. One of our solutions in particular has cosmological symmetry, i.e. it is spatially homogeneous and isotropic, and it is additionally conformally flat, with conformal factor depending only on the timelike coordinate. We show that this conformal factor can be interpreted as the scale factor, we compute it as a function of cosmological time, and we show that it corresponds to a linearly expanding (or contracting) Finsler universe.

## I. INTRODUCTION

The interest in Finsler-geometric modified theories of gravity has picked up in recent years, and rightly so. It has become clear that small, natural modifications in the classic axiomatic approach by Ehlers, Pirani and Schild (EPS) [1] naturally lead to Finsler spacetime geometry; [2–4]; it has also become clear that modified dispersion relations (MDRs), usually discussed in the context of quantum gravity phenomenology [5] generically induce a Finsler geometry on spacetime [6–9] and it has been conjectured that Finsler spacetime geometry describes the gravitational field of a kinetic gas more accurately compared to its usual treatment in the Einstein-Vlasov system, [10, 11]. These results, to name just a few, show clearly that in certain situations, for instance in Planckian regimes or for certain type of matter, it is to be expected that Finsler geometry should be the proper way to model spacetime. Beside the applications in gravitational physics, see also [12], there are several other instances in physics, such as the description of the propagation of waves in media, where Finsler geometry seems to be the appropriate tool [13].

Various physically subclasses of Finsler spacetimes can be distinguished, and two particularly important ones are the class of Berwald spacetimes and the class of Landsberg spacetimes that may be thought of as

incrementally non-(pseudo-)Riemannian, respectively (precise definitions will follow later). Every Berwald spacetime is also Landsberg, but whether or not the opposite is true has been a long standing open question in Finsler geometry. In fact, Matsumoto has stated in 2003 that this question represents the next frontier of Finsler geometry [14], and as a token of their elusivity, Bao [14] has called these non-Berwaldian Landsberg spaces ‘[...] unicorns, by analogy with those mythical single-horned horse-like creatures for which no confirmed sighting is available.’ Since 2006 some examples of unicorns have been obtained by Asanov [15], Shen [16] and Elgendi [17] by relaxing the definition of a Finsler space. Even such examples of so-called  $y$ -local unicorns are still exceedingly rare.

Here we present a new family of exact solutions to Pfeifer and Wohlfarth’s Finslerian extension of Einstein’s field equations [18, 19] which is precisely such a unicorn. It falls into one of the classes introduced by Elgendi. Our solutions extend the very short list of known exact solutions in Finsler gravity. Indeed, to the best of our knowledge the only ones currently known in the literature are the (m-Kropina type) Finsler pp-waves [20] and their generalization as Very General Relativity (VGR) spacetimes [21], the Randers pp-waves [22], and the pp-waves of general  $(\alpha, \beta)$ -metric type [23].

Interestingly we find that these solutions have a physically viable light cone structure, even though in some cases the signature is not Lorentzian but positive definite. In fact, the Finslerian light cone turns out to be equivalent to that of the flat Minkowski metric. Furthermore we find a natural cosmological interpretation of one

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of our solutions and a promising analogy with classical FLRW cosmology. In particular, our solution has cosmological symmetry, i.e. it is spatially homogeneous and isotropic, and it is additionally conformally flat, with conformal factor depending only on the timelike coordinate. We show that this conformal factor can be interpreted as the scale factor, we compute the scale factor as a function of cosmological time, and we show that it corresponds to a linearly expanding (or contracting) Finslerian universe.

## II. FINSLER GEOMETRY

Before we recall the basic notions of Finsler geometry below we first introduce some notation. Given a (spacetime) manifold  $M$ , which we assume to be 4-dimensional, and given some coordinates on  $M$ , we will always consider the natural induced coordinates on its tangent bundle  $TM$ . More precisely, given a coordinate chart  $(U, x)$  on  $M$ ,  $U \subset M$ , we obtain a coordinate chart  $(TU, (x, y))$  on  $TM$ ,  $TU \subset TM$  of  $TM$ , where a point  $Y = y^\mu \partial_\mu \in TM$  is labeled by  $(x, y)$ . By a slight abuse of notation we will generally identify any point in  $M$  with its expression in coordinates, and similarly for points in  $TM$ , i.e.  $Y = y$ . We denote the coordinate basis vectors of the tangent spaces of  $TM$  by  $\partial_\mu = \partial/\partial x^\mu$  and  $\bar{\partial}_\mu = \partial/\partial y^\mu$ , where  $\mu = 0, \dots, 3$ .

A Finsler space is a smooth manifold  $M$  endowed with a Finsler metric, i.e. a smooth map  $F : TM \setminus 0 \rightarrow \mathbb{R}$  such that

- $F$  is (positively) homogeneous of degree one with respect to  $y$ :

$$F(x, \lambda y) = \lambda F(x, y), \quad \forall \lambda > 0; \quad (1)$$

- the *fundamental tensor*

$$g_{\mu\nu} = \bar{\partial}_\mu \bar{\partial}_\nu \left( \frac{1}{2} F^2 \right) \quad (2)$$

is non-degenerate.

The fundamental tensor  $g_{\mu\nu}$  depends generally on both  $x$  and  $y$ . When  $F$  is quadratic in  $y$ , or equivalently when  $g_{\mu\nu}$  depends only on  $x$ , then  $g_{\mu\nu}$  is a (pseudo-)Riemannian metric and the theory reduces to (pseudo-)Riemannian geometry.

In order to describe spacetime geometry, one usually demands that the signature of  $g_{\mu\nu}$  be Lorentzian, at least in some conic open subset of  $TM$ , which one might hope to identify with the cone of timelike directions. Moreover, in applications one very often encounters Finsler structures<sup>1</sup> that are only properly defined (smooth, non-

degenerate) on a subset of  $TM \setminus 0$ . Such Finsler metrics are sometimes referred to as  $y$ -local, as opposed to  $y$ -global [14]. In particular, the unicorn solution that we will present here is of  $y$ -local type.

### A. The nonlinear connection and geodesic spray

The *Cartan non-linear connection* is the unique homogeneous (in general non-linear) Ehresmann connection on  $TM$  that is smooth on  $TM \setminus \{0\}$ , torsion-free and compatible with  $F$ . It may therefore be viewed as a generalization of the Levi-Civita connection. For details we refer e.g. to [28]. Its connection coefficients are given by

$$N_\nu^\mu = \frac{1}{4} \bar{\partial}_\nu \left( g^{\mu\rho} (y^\sigma \partial_\sigma \bar{\partial}_\rho F^2 - \partial_\rho F^2) \right). \quad (3)$$

The nonlinear connection induces the horizontal derivatives

$$\delta_\mu = \partial_\mu - N_\mu^\nu \bar{\partial}_\nu, \quad (4)$$

that, together with the  $\bar{\partial}_\mu$ , span each tangent space  $T_{(x, \dot{x})} TM$ . The (geodesic) spray coefficients can then be defined as

$$G^\mu \equiv N_\nu^\mu y^\nu = \frac{1}{2} g^{\mu\rho} (y^\sigma \partial_\sigma \bar{\partial}_\rho L - \partial_\rho L), \quad (5)$$

where the second equality follows from Euler's theorem for homogeneous functions. It immediately follows that we also have  $N_\nu^\mu = \frac{1}{2} \bar{\partial}_\nu G^\mu$ . The importance of the spray coefficients comes from the fact that the geodesics of  $F$  are given by

$$\ddot{x}^\mu + G^\mu(x, \dot{x}) = 0, \quad (6)$$

which coincidentally is also the autoparallel equation of the nonlinear connection  $N_\nu^\mu$ .

The curvature of the nonlinear connection is defined via  $R^\rho{}_{\mu\nu} \bar{\partial}_\rho = -[\delta_\mu, \delta_\nu]$ , which implies that

$$R^\rho{}_{\mu\nu} = \delta_\mu N_\nu^\rho - \delta_\nu N_\mu^\rho. \quad (7)$$

From the nonlinear curvature one may define the Finsler Ricci scalar and Ricci tensor as follows

$$\text{Ric} = R^\rho{}_{\rho\mu} y^\mu, \quad R_{\mu\nu} = \frac{1}{2} \bar{\partial}_\mu \bar{\partial}_\nu \text{Ric}. \quad (8)$$

A Finsler space is said to be *Ricci-flat* if  $\text{Ric} = 0$ , or equivalently,  $R_{\mu\nu} = 0$ . We remark that the Finsler Ricci scalar is not to be confused with the scalar curvature usually defined in (pseudo-)Riemannian geometry as  $R = g^{\mu\nu} R_{\mu\nu}$ , also sometimes called the Ricci scalar.

### B. The Chern-Rund connection

In addition to the canonical nonlinear connection, various canonical linear connections can be introduced. However, the price one has to pay for linearity is that the

<sup>1</sup> In the literature one finds various other, more stringent, definitions of Finsler spacetimes, going back to the original definition by Beem [24]. They vary in their precise technical details, depending on the scope of the application, see e.g. [25–27] and references therein.

linear connections do not in general live on the vector bundle  $TM$  but rather on its pull-back  $\pi^*TM$  by the canonical projection  $\pi : TM \rightarrow M$ . The pull-back bundle  $\pi^*TM$  is considered as a vector bundle over  $TM \setminus 0$  and sections of this vector bundle may be thought of simply as vector fields on  $M$  with a dependence on both  $x$  and  $y \neq 0$ . Since the manifold  $TM \setminus 0$  has dimension  $2n$ , we get in general two sets of linear connection coefficients, namely

$$\nabla_{\delta_\mu} \partial_\nu = \Gamma_{\nu\mu}^\rho \partial_\rho, \quad \nabla_{\bar{\delta}_\mu} \partial_\nu = \bar{\Gamma}_{\nu\mu}^\rho \partial_\rho. \quad (9)$$

The Chern-Rund connection is the unique linear connection on  $\pi^*TM$  that is torsion-free and *almost* metric compatible. For details we refer e.g. to [29]. These conditions imply that  $\bar{\Gamma}_{\mu\nu}^\rho = 0$  and

$$\Gamma_{\mu\nu}^\rho = \frac{1}{2} g^{\rho\sigma} (\delta_\mu g_{\sigma\nu} + \delta_\nu g_{\mu\sigma} - \delta_\sigma g_{\mu\nu}). \quad (10)$$

Notice the similarity to the formula for the Levi-Civita Christoffel symbols of a (pseudo-)Riemannian metric. From this it is immediately clear that the Chern-Rund connection reduces to the Levi-Civita connection when  $F$  is (pseudo-)Riemannian.

### C. Berwald and Landsberg spaces

Next we introduce two important classes of Finsler spaces: Berwald spaces and Landsberg spaces. First, if the spray is quadratic in  $y$ , i.e.  $\bar{\delta}_\mu \bar{\delta}_\nu \bar{\delta}_\sigma G^\rho = 0$  then  $F$  is said to be of *Berwald* type. What this means geometrically is that the Chern connection may be understood as an affine connection on  $M$ , i.e. equivalently, a space is Berwald if and only if the connection coefficients  $\Gamma_{\mu\nu}^\rho$  of the Chern connection depend only on  $x$ .

And second, introducing the Landsberg curvature

$$S_{\mu\nu\sigma} = -\frac{1}{4} y_\rho \bar{\delta}_\mu \bar{\delta}_\nu \bar{\delta}_\sigma G^\rho, \quad (11)$$

and the mean Landsberg curvature  $S_\sigma = g^{\mu\nu} S_{\sigma\mu\nu}$ , we say that a space is (weakly) *Landsberg* if the (mean) Landsberg curvature vanishes identically. The geometrical significance of the Landsberg tensor is somewhat more difficult to state in simple terms without introducing more machinery, so instead we refer e.g. to [14].

It is immediately obvious from the definitions that any Berwald space is a Landsberg space. Also, a (pseudo-)Riemannian space is always Berwald, hence in particular any (pseudo-)Riemannian space is Landsberg.

### D. Unicorns in Finsler geometry

As observed above, we have the following inclusions:

$$(\text{pseudo-})\text{Riemannian} \subset \text{Berwald} \subset \text{Landsberg}.$$

It has been a long standing open question whether the last inclusion is strict. Do there exist Landsberg space that are not Berwald? In the  $y$ -global case the answer is unknown. For  $y$ -local spaces some examples are known, but these are exceedingly rare. As such, non-Berwaldian Landsberg spaces are referred to as *unicorns* [14]. We recommend [14, 30] for reviews on the unicorn problem.

The first unicorns were found by Asanov [15] in 2006 and his results were generalized by Shen [16] a few years later. These were the only known examples of unicorns until Elgendi very recently provided some additional examples of unicorns [17]. One of the families of unicorns introduced by Elgendi will be central in this work.

## III. THE FINSLERIAN FIELD EQUATIONS

Although various proposals for Finslerian field equations in vacuum can be found in the literature [18, 31–40], it seems fair to say that Pfeifer and Wohlfarth's field equation [18] has the most robust foundation. It is obtained as the Euler-Lagrange equation of the natural Finsler generalization of the Einstein Hilbert action [18, 19], and furthermore it has been shown recently that the equation is the *variational completion* of Rutz's equation [31],  $\text{Ric} = 0$ . The latter is arguably the simplest and cleanest proposal, and well physically motivated, but it cannot be obtained by extremizing an action functional, complicating the coupling of the theory to matter. For reference, Einstein's vacuum equation in the form  $R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R = 0$  is also precisely the variational completion of the simpler equation  $R_{\mu\nu} = 0$  [41]. While in the GR case the completed equation happens to be equivalent to the former, this is not true any longer in the Finsler setting.

Pfeifer and Wohlfarth's field equation in vacuum reads

$$\begin{aligned} \text{Ric} - \frac{L}{3} g^{\mu\nu} R_{\mu\nu} \\ - \frac{L}{3} g^{\mu\nu} \left( \bar{\delta}_\mu \dot{S}_\nu - S_\mu S_\nu + \nabla_{\delta_\mu} S_\nu \right) = 0, \end{aligned} \quad (12)$$

where  $\dot{S}_\nu \equiv y^\rho \nabla_{\delta_\rho} S_\nu$ . For (pseudo-)Riemannian metrics, (12) reduces to Einstein's field equation in vacuum. From the general expression (12) it becomes immediately apparent that for weakly Landsberg spaces, characterized by the defining property that  $S_i = 0$ , the field equation in vacuum attains the relatively simple form

$$\text{Ric} - \frac{L}{3} g^{\mu\nu} R_{\mu\nu} = 0. \quad (13)$$

Recalling the definition (8) of  $R_{\mu\nu}$  we have the following immediate result:

**Proposition 1.** *Any Ricci-flat,  $\text{Ric} = 0$ , weakly Landsberg space is a solution to the field equations (12)*

In other words, any weakly Landsberg solution to the Rutz equation is automatically a solution to (12).

## IV. AN EXACT UNICORN SOLUTION TO FINSLER GRAVITY

### A. Elgendi's class of unicorns

Elgendi recently introduced a class of unicorns [17] with Finsler metric given by

$$F = \left( a\beta + \sqrt{\alpha^2 - \beta^2} \right) e^{\frac{a\beta}{a\beta + \sqrt{\alpha^2 - \beta^2}}}, \quad (14)$$

in terms of a real, nonvanishing constant  $a$  and

$$\alpha = f(x^0)\sqrt{(y^0)^2 + \phi(\hat{y})}, \quad \beta = f(x^0)y^0, \quad (15)$$

where  $f$  is a real-valued function and  $\phi(\hat{y}) = \phi_{ij}\hat{y}^i\hat{y}^j = \phi_{ij}y^iy^j$  is a non-degenerate quadratic form on the space spanned by  $\hat{y} = (y^1, y^2, y^3)$ , with constant, symmetric coefficients  $\phi_{ij}$ . Here and in what follows, indices  $i, j, \dots$  will run over 1, 2, 3, whereas greek induces  $\mu, \nu, \dots$  will run over 0, 1, 2, 3. From a gravitational physics perspective, the only degree of freedom of these Finsler functions is the function  $f(x^0)$ . The geodesic spray of  $F$  is given by

$$G^0 = \left( \frac{2f(x^0)^2(y^0)^2 - \alpha^2}{f(x^0)^2} + \frac{a^2 - 1}{a^2} \frac{\alpha^2 - \beta^2}{f(x^0)^2} \right) \frac{f'(x^0)}{f(x^0)} \quad (16)$$

$$G^i = Py^i, \quad (17)$$

where

$$P = 2 \left( y^0 + \frac{1}{af(x^0)} \sqrt{\alpha^2 - \beta^2} \right) \frac{f'(x^0)}{f(x^0)}, \quad (18)$$

and the Landsberg tensor vanishes identically. Note that our  $G^k$  is twice the  $G^k$  in Elgendi's paper [17], due to a difference in convention. Explicitly then, Elgendi's unicorns have the form

$$F = f(x^0) \left( y^0 + \sqrt{\phi(\hat{y})} \right) e^{\frac{y^0}{y^0 + \sqrt{\phi(\hat{y})}}}. \quad (19)$$

where we have absorbed the constant  $a$  into a redefinition of  $x^0$ . For our purposes we will modify this expression slightly, though.

### B. The modified unicorn metric

The expression (19) defining the unicorn metric is only well-defined whenever  $\phi(\hat{y}) \geq 0$ . If  $\phi$  is positive definite, this is necessarily the case, but in other signatures this is not always true. In order to extend the domain of definition of  $F$ , an obvious first approach would naturally be to replace  $\phi$  by its absolute value,  $|\phi|$ , i.e.

$$F = f(x^0) \left( y^0 + \sqrt{|\phi(\hat{y})|} \right) e^{\frac{y^0}{y^0 + \sqrt{|\phi(\hat{y})|}}}. \quad (20)$$

From the physical point of view this is still not completely satisfactory, however. This can be seen by considering the light cone corresponding to such a Finsler metric  $F$ , given by the set of vectors for which  $F = 0$ . Indeed, the light cone would be given by those vectors satisfying  $y^0 = -\sqrt{|\phi|}$ , which would imply that the light cone is contained entirely within the half space  $y^0 < 0$ , which does not seem very realistic. To obtain a viable light cone structure we consider a modified unicorn metric, inspired by the construction of modified Randers metrics in [23].

Thus, our starting point will be the following *modified unicorn metric*:

$$F = f(x^0) \left( |y^0| + \text{sgn}(\phi)\sqrt{|\phi|} \right) e^{\frac{|y^0|}{|y^0| + \text{sgn}(\phi)\sqrt{|\phi|}}}. \quad (21)$$

Below we'll show that such a metric is still of the unicorn type, and indeed defines a physically reasonable cone structure, under some conditions on  $\phi$ . After that, we will determine the signature of the fundamental tensor, which will turn out to depend on the signature of  $\phi$ , and finally, in section IV C we determine the free function  $f(x^0)$  in (21) by application of the Finsler gravity equation. Afterwards we discuss the physical interpretation and conclude.

#### 1. Cone structure

First we observe that, regardless of the exact form or signature of  $\phi$ , our modified unicorn metrics have a light cone structure that is equivalent to that of a pseudo-Riemannian metric.

**Proposition 2.** *The light cone of the modified unicorn metric (21) is given by*

$$(y^0)^2 + \phi = 0. \quad (22)$$

*Proof.* The result follows from the following sequence of equivalences.

$$F = 0 \Leftrightarrow \text{sgn}(\phi)\sqrt{|\phi|} + |y^0| = 0 \quad (23)$$

$$\Leftrightarrow \text{sgn}(\phi)\sqrt{|\phi|} = -|y^0| \quad (24)$$

$$\Leftrightarrow |\phi| = (y^0)^2 \quad \text{and} \quad \phi < 0 \quad (25)$$

$$\Leftrightarrow \phi = -(y^0)^2 \quad (26)$$

$$\Leftrightarrow \phi + (y^0)^2 = 0. \quad (27)$$

□

Depending on the signature of  $\phi$ , we can make a more precise statement.

**Proposition 3.** *Let  $F$  be the modified unicorn metric (21) corresponding to some non-degenerate quadratic form  $\phi_{ij}$ . Then the following holds:*

- For  $\phi_{ij}$  of signature  $(+, +, -)$  the null structure of  $F$  is identical to the Minkowski metric spacetime lightcone structure of signature  $(+, +, +, -)$ .
- If  $\phi_{ij}$  is negative definite, i.e. of signature  $(-, -, -)$  then the light cone of  $F$  is identical to the Minkowski metric spacetime lightcone of signature  $(+, -, -, -)$ .
- If  $\phi_{ij}$  is positive definite, i.e. of signature  $(+, +, +)$  then the light cone of  $F$  is given by  $y^\mu = 0$ .
- For  $\phi_{ij}$  of signature  $(-, -, x)$  the null structure of  $F$  is identical to the one of a pseudo-Riemannian metric manifold with signature  $(+, -, -, +)$ .

This singles out the  $(+, +, -)$  and  $(-, -, -)$  signatures of  $\phi_{ij}$  as the ones that are physically reasonable. In the first case it leads to the interpretation of the coordinate  $x^3$  as timelike coordinate, while in the second case  $x^0$  would be the timelike coordinate.

## 2. Signature of the fundamental tensor

Next we investigate the signature of our modified unicorn metrics.

**Proposition 4.** Consider a modified unicorn metric  $F$  as in (21) and let  $\mathcal{S}(\phi)$  be the set of all  $\hat{y} = (y^1, y^2, y^3)$  that are  $\phi$ -spacelike and  $\mathcal{T}(\phi)$  the set of all  $\hat{y}$  that are  $\phi$ -timelike.

- If  $\phi$  is positive definite or negative definite then  $g_{\mu\nu}(x, y)$  is positive definite on its entire domain of definition.
- If  $\phi$  is Lorentzian then  $g_{\mu\nu}(x, y)$  is of Lorentzian signature  $(+, +, +, -)$  for all  $y \in \mathbb{R} \times \mathcal{S}(\phi)$  and  $g_{\mu\nu}(x, y)$  is of signature  $(+, -, -, +)$  for all  $y \in \mathbb{R} \times \mathcal{T}(\phi)$ .

Before we give the proof, let us make some interesting observations. Consider the two physically reasonable scenario's we identified below Prop. 3 as a result of their good cone structure, i.e.  $\phi$  having  $(+, +, -)$  or  $(-, -, -)$  signature. In both cases the light cone is equivalent to that of Minkowski space. Surprisingly, inside the interior of this cone it is easy to see from the previous proposition, that the signature of  $g$  is not Lorentzian. Indeed, if  $\phi$  has signature  $(+, +, -)$ , then  $g$  has signature  $(+, -, -, +)$  inside the cone, while if  $\phi$  is negative definite then  $g$  is positive definite inside the cone. In the later case this does not contradict the existence of the lightcone structure due to the reduced smoothness of the Finsler function  $F$  (21).

In many definitions of Finsler spacetimes [25, 27], one requires that there exists a cone inside which the fundamental tensor has Lorentzian signature, in order to guarantee the existence of a physical cone structure. Here,

however, we found that there exist Finsler geometries, which do have a satisfactory cone structure, without this property. This is an interesting new observation about Finsler geometry in its own right: apparently even in positive definite signature, a cone structure may arise due to irregularities (i.e. non-smoothness) of the Finsler metric. It is therefore not immediately obvious whether the positive definite signature poses a fundamental problem in the application of the unicorn metrics in spacetime physics or not. In the context of standard static Finsler spacetimes, a lack of smoothness of the Finsler function inside the cone was also considered acceptable [26, 42].

*Proof of Prop. 4.* We may choose coordinates such that  $\phi = \phi(\hat{y}) = \varepsilon_1(y^1)^2 + \varepsilon_2(y^2)^2 + \varepsilon_3(y^3)^2$ . Since the spacetime dimension is fixed to 4, wherever  $F$  is sufficiently differentiable the calculation of the determinant of the fundamental tensor is in principle a straightforward exercise. It is given by

$$\det g = \text{sgn}(\phi(\hat{y}))\varepsilon_1\varepsilon_2\varepsilon_3f(x^0)^8 \exp\left(\frac{8|y^0|}{|y^0| + \text{sgn}(\phi)\sqrt{|\phi|}}\right). \quad (28)$$

The determinant already gives us a pretty good idea of what the possible signature of  $g_{\mu\nu}$  can be. In particular, since  $g_{\mu\nu}$  is a four-dimensional matrix, it has Lorentzian signature, either of type  $(+, -, -, -)$  or  $(-, +, +, +)$ , if and only if its determinant is negative.

- Assume  $\phi_{ij}$  is positive definite, then all  $\varepsilon_i$  and  $\text{sgn}(\phi(\hat{y}))$  are positive, and hence  $\det g$  is.
- Assume  $\phi_{ij}$  is negative definite, then all  $\varepsilon_i$  and  $\text{sgn}(\phi(\hat{y}))$  are negative, and hence  $\det g$  is positive.
- Assume  $\phi_{ij}$  is of Lorentzian signature  $(+, +, -)$ , then  $\det g$  is negative whenever  $\text{sgn}(\phi(\hat{y})) > 0$ , i.e. on  $\mathcal{S}(\phi)$ , and  $\det g$  is positive whenever  $\text{sgn}(\phi(\hat{y})) < 0$ , i.e. on  $\mathcal{T}(\phi)$ .
- Assume  $\phi_{ij}$  is of Lorentzian signature  $(+, +, -)$ , then  $\det g$  is negative whenever  $\text{sgn}(\phi(\hat{y})) < 0$ , i.e. on  $\mathcal{S}(\phi)$ , and  $\det g$  is positive whenever  $\text{sgn}(\phi(\hat{y})) > 0$ , i.e. on  $\mathcal{T}(\phi)$ .

This already shows that  $g_{\mu\nu}$  is Lorentzian if and only if  $\phi$  is Lorentzian and  $y \in \mathbb{R} \times \mathcal{S}(\phi)$ . But the sign of the determinant does not suffice to determine whether this signature is mostly plus or mostly minus. Similarly it does not tell us much about the signature of  $g_{\mu\nu}$  when  $\phi$  is positive or negative definite. In order to found, we'll distinguish the following cases.

### Case 1: $\phi$ Lorentzian and $y \in \mathbb{R} \times \mathcal{S}(\phi)$

We first consider the case that  $\phi$  is Lorentzian. Without loss of generality (wlog) we set  $\phi(\hat{y}) = \varepsilon(y^1)^2 + \varepsilon(y^2)^2 - \varepsilon(y^3)^2$ , where  $\varepsilon = \pm 1$ . The choice of the sign  $\varepsilon$  selects if we are in case c) or d) from above.

Now note that given a vector  $y \in T_x M$  which is  $\phi$ -spacelike, it follows from the symmetries of the Finsler metric and in particular from the 3-dimensional Lorentz symmetry of  $\phi$  that for we may always change coordinates, without changing the form of  $\phi$  (and  $F$ ), such that  $y^2 = y^3 = 0$ .

For any choice of epsilon, by direct calculation we find, using that  $\epsilon^2 = 1$  and  $|\epsilon| = 1$ , that  $g_{\mu\nu}$  is of the form

$$g_{\mu\nu} = e^{\frac{2|y^0|}{|y^0| + \epsilon|y^1|}} f(x^0)^2 \begin{pmatrix} M & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & -1 \end{pmatrix}, \quad (29)$$

where  $M$  is an ( $\epsilon$ -dependent) positive definite  $2 \times 2$  matrix<sup>2</sup>. Hence we conclude that  $g_{\mu\nu}$  is of the mostly plus type  $(+, +, +, -)$ .

### Case 2: $\phi$ Lorentzian and $y \in \mathbb{R} \times \mathcal{T}(\phi)$

In this case we may wlog choose coordinates such that  $\phi(\hat{y}) = -\epsilon(y^1)^2 + \epsilon(y^2)^2 + \epsilon(y^3)^2$ , where  $\epsilon = \pm 1$ , and such that  $y^2 = y^3 = 0$ . Again by direct calculation we find that  $g_{\mu\nu}$  is of the form

$$g_{\mu\nu} = e^{\frac{2|y^0|}{|y^0| + \epsilon|y^1|}} f(x^0)^2 \begin{pmatrix} M & 0 & 0 \\ 0 & -1 & 0 \\ 0 & 0 & -1 \end{pmatrix}, \quad (30)$$

where  $M$  is a positive definite  $2 \times 2$  matrix. Hence we conclude that  $g_{\mu\nu}$  is in this case of signature  $(+, -, -, +)$ .

### Case 3: $\phi$ positive or negative definite

In this case we may wlog choose coordinates such that  $\phi(\hat{y}) = \epsilon(y^1)^2 + \epsilon(y^2)^2 + \epsilon(y^3)^2$ , where  $\epsilon = \pm 1$ , and such that, for any given  $y \in T_x M$ , we have  $y^2 = y^3 = 0$ . Again by direct calculation we find that  $g_{\mu\nu}$  is of the form

$$g_{\mu\nu} = e^{\frac{2|y^0|}{|y^0| + \epsilon|y^1|}} f(x^0)^2 \begin{pmatrix} M & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & 1 \end{pmatrix}, \quad (31)$$

where  $M$  is a positive definite  $2 \times 2$  matrix. Hence we conclude that  $g_{\mu\nu}$  is positive definite.  $\square$

## C. Solving the Finsler gravity equation

Next, we seek to determine the form of the free function  $f(x^0)$  of (21) from the Finsler gravity equations (13).

The geodesic spray of  $F$  is given explicitly by

$$G^0 = ((y^0)^2 - |\phi|) \frac{f'(x^0)}{f(x^0)} \quad (32)$$

$$G^i = P y^i, \quad i = 1, 2, 3 \quad (33)$$

where

$$P = 2 \left( |y^0| + \text{sgn}(\phi) \sqrt{|\phi|} \right) \text{sgn}(y^0) \frac{f'(x^0)}{f(x^0)} \quad (34)$$

**Proposition 5.**  *$F$  is Ricci-flat if and only if  $f$  has the form  $f(x^0) = c_1 \exp(c_2 x^0)$ .*

*Proof.* By definition, and using homogeneity and the fact that  $N_\nu^\mu = \frac{1}{2} \bar{\partial}_\nu G^\mu$ , we have

$$\text{Ric} = R^\mu{}_{\nu\mu\nu} y^\nu = (\delta_\mu N_\nu^\mu - \delta_\nu N_\mu^\mu) y^\nu \quad (35)$$

$$= y^\nu ((\partial_\mu - N_\mu^\rho \bar{\partial}_\rho) N_\nu^\mu - (\partial_\nu - N_\nu^\rho \bar{\partial}_\rho) N_\mu^\mu) \quad (36)$$

$$= y^\nu \partial_\mu N_\nu^\mu - y^\nu \partial_\nu N_\mu^\mu \quad (37)$$

$$- y^\nu N_\mu^\rho \bar{\partial}_\rho N_\nu^\mu + y^\nu N_\nu^\rho \bar{\partial}_\rho N_\mu^\mu \quad (38)$$

$$= \frac{1}{2} (y^\nu \partial_\mu \bar{\partial}_\nu G^\mu - y^\nu \partial_\nu \bar{\partial}_\mu G^\mu) \quad (39)$$

$$- \frac{1}{4} (y^\nu \bar{\partial}_\mu G^\rho \bar{\partial}_\rho \bar{\partial}_\nu G^\mu - y^\nu \bar{\partial}_\nu G^\rho \bar{\partial}_\rho \bar{\partial}_\mu G^\mu) \quad (40)$$

$$= \partial_\mu G^\mu - \frac{1}{2} y^\nu \partial_\nu \bar{\partial}_\mu G^\mu \quad (41)$$

$$- \frac{1}{4} (\bar{\partial}_\mu G^\rho \bar{\partial}_\rho G^\mu - 2G^\rho \bar{\partial}_\rho \bar{\partial}_\mu G^\mu). \quad (42)$$

Using the identities

$$\begin{aligned} \bar{\partial}_0 P &= 2f'/f, & \bar{\partial}_0 G^0 &= 2y^0 f'/f, \\ \bar{\partial}_0^2 G^0 &= 2f''/f, & \bar{\partial}_0 \bar{\partial}_i G^0 &= \bar{\partial}_0 \bar{\partial}_i P = \bar{\partial}_0^2 P = 0, \\ y^i \bar{\partial}_i G^0 &= -2|\phi| f'/f, & y^i \bar{\partial}_i P &= 2\text{sgn}(\phi y^0) \sqrt{|\phi|} f'/f, \end{aligned} \quad (43)$$

one finds after some slightly tedious manipulations that the last two terms in the expression for the Ricci tensor can both be expressed as

$$\bar{\partial}_\mu G^\rho \bar{\partial}_\rho G^\mu = 2G^\rho \bar{\partial}_\rho \bar{\partial}_\mu G^\mu = nP^2, \quad (44)$$

where  $n = \dim M = 4$  in our case. Hence these terms cancel each other out precisely. Denoting  $G^0 = \bar{G}^0 f'/f$  and  $P = \bar{P} f'/f$ , so that  $\bar{G}^0$  and  $\bar{P}$  do not depend on  $x^\mu$ , and using the fact that  $\bar{\partial}_\mu G^\mu = nP$ , one finds furthermore that

$$\partial_\mu G^\mu = (\partial_0^2 \log |f|) \bar{G}^0, \quad (45)$$

$$y^\nu \partial_\nu \bar{\partial}_\mu G^\mu = n y^0 (\partial_0^2 \log |f|) \bar{P}. \quad (46)$$

Consequently we have

$$\text{Ric} = \partial_\mu G^\mu - \frac{1}{2} y^\nu \partial_\nu \bar{\partial}_\mu G^\mu \quad (47)$$

$$= +(\partial_0^2 \log |f|) \quad (48)$$

$$\times \left( (1-n)(y^0)^2 - n|y^0| \text{sgn}(\phi) \sqrt{|\phi|} - |\phi| \right), \quad (49)$$

<sup>2</sup>  $M$  must be either positive definite or negative definite, since we have already shown that  $g_{ij}$  has Lorentzian signature. The fact that  $M_{ab} v^a v^b = ((y^0)^2 + (y^1)^2) / (|y^0| + \epsilon|y^1|)^2 > 0$  for  $v^a = (0, 1)$  thus shows that  $M$  is positive definite.

which in dimension  $n = 4$  reduces to

$$\text{Ric} = -(\partial_0^2 \log |f|) \quad (50)$$

$$\times \left( 3(y^0)^2 + 4|y^0| \text{sgn}(\phi) \sqrt{|\phi|} + |\phi| \right). \quad (51)$$

If  $\text{Ric} = 0$  then we must in particular have

$$0 = \bar{\partial}_0^2 \text{Ric} = -6(\partial_0^2 \log |f|). \quad (52)$$

It thus follows that  $\text{Ric} = 0$  if and only if  $\partial_0^2 \log |f| = 0$ , the general solution to which is given by the stated form of  $f$ .  $\square$

This shows that the following family of Finsler metrics are exact vacuum solutions to Pfeifer and Wohlfarth's field equations in Finsler gravity:

$$F = c_1 e^{c_2 x^0} \left( |y^0| + \text{sgn}(\phi) \sqrt{|\phi|} \right) \exp \left( \frac{|y^0|}{|y^0| + \text{sgn}(\phi) \sqrt{|\phi|}} \right), \quad (53)$$

where  $\phi = \phi(\hat{y}) = \phi_{ij} y^i y^j$ , with  $\phi_{ij}$  being a three-dimensional non-degenerate, symmetric bilinear form with constant coefficients and signature  $(+, +, -)$  or  $(-, -, -)$ .

In fact it turns out that *any* solution of the type (21) must have this form.

**Proposition 6.** *A Finsler metric of the form (21) is a solution to the Finslerian field equations in vacuum, Eq. (13), if and only if it can be written locally as (53).*

Before we give a sketch of the proof, it is important to point out the exact meaning of the word 'locally' in the proposition, as it will have essential consequences for the physical viability of such solutions. Of course, the word applies first and foremost to the  $x$ -coordinates in the usual sense, but it also applies to the tangent space coordinates  $y^i$ : there is *a priori* no reason why one couldn't pick, at some point  $x^\mu$ , say, a certain  $\phi_{ij}$  in a certain subset of  $T_x M$  and a different  $\phi_{ij}$  in a different subset of  $T_x M$ . Of course this could in general result in not having smoothness across the interface of the two regions, but this does not necessarily have to be a problem, unless one sets very strict smoothness requirements.

For consistency the different parts of  $TM$  must satisfy some conditions. It is most natural to require that any such part be a conic subbundle, i.e. an open conic subset with non-empty fiber at each point in  $M$ .

*Proof.* In four spacetime dimensions the proof is straightforward and most easily performed in convenient coordinates in which  $\phi$  is diagonal with all nonvanishing entries equal to  $+1$  or  $-1$ . From (21) one can directly compute  $g_{\mu\nu}$  and then its inverse  $g^{\mu\nu}$ . From (51) together with (8) one can immediately compute  $R_{\mu\nu}$ . We omit the intermediate expressions because they are somewhat lengthy,

but plugging all of this into the field equation (13) leads to

$$\frac{-\partial_0^2 \log |f|}{3\sqrt{|\phi|}} \left( -4 \text{sgn}(\phi) |y^0|^3 - 5y^0 \sqrt{|\phi|} + |\phi|^{3/2} \right) = 0. \quad (54)$$

Clearly this equation can only be satisfied for all  $y^\mu$  in an open set where the fundamental tensor has Lorentzian signature if  $\partial_0^2 \log |f| = 0$ , in which case (the last sentence in the proof of) Prop. 5 shows that  $F$  is in fact Ricci-flat and therefore must have the form (53).  $\square$

With this we found an exact solution to the Finsler gravity equations (53), starting from a generalised version of Elgendi's unicorns (21).

#### D. Physical interpretation: A linearly expanding universe

Having analyzed the mathematical properties of the unicorn Finsler spacetimes (21), and found the exact unicorn vacuum solution (53) of the Finsler gravity vacuum equations (13), we now turn to the physical interpretation of this solution. We find that these Finsler gravity vacuum solutions yield a vacuum cosmology, with linear time dependence of the scale factor.

To reach this conclusion we highlight the following properties of the the unicorn Finsler spacetimes (21):

- Conformal flatness, with a conformal factor that is only spacetime dependent

$$F(x, y) = f(x) F_0(y). \quad (55)$$

- Cosmological symmetry, for the case when  $\phi_{ij}$  has signature  $(-, -, -)$ , since then, by introducing spatial spherical coordinates  $(r, \theta, \phi)$ , we can write

$$F(x, y) = F(x^0, y^0, w), \quad (56)$$

$$w^2 = (y^1)^2 + (y^2)^2 + (y^3)^2 \quad (57)$$

$$= (y^r)^2 + r^2 ((y^\theta)^2 + \sin^2 \theta (y^\phi)^2), \quad (58)$$

which is precisely of the form of a spatial flat homogeneous and isotropic Finsler geometry [43]. This construction does not work for  $\phi_{ij}$  with signature  $(+, +, -)$ , since then  $y^3$ , and not  $y^1$  would be the timelike direction.

Combining these observations, we find that the unicorn Finsler spacetimes (21) are of the form

$$F(x^0, y^0, w) = f(x^0) F_0(y^0, w). \quad (59)$$

This from reminds immediately at classical flat FLRW spacetimes in conformal time  $\eta$ , which we identify here

with the  $x^0$  coordinate. When written in the language of Finsler geometry they are of the form (59) with

$$F_{0FLRW} = \sqrt{-(y^0)^2 + (y^1)^2(y^2)^2 + (y^3)^2}. \quad (60)$$

A redefinition of the time coordinate via  $\frac{\partial \tilde{x}^0}{\partial x^0} = f(x^0)$ , which implies that  $\tilde{y}^0 = f(x^0)y^0$ , then leads to the standard form of flat FLRW geometry

$$F = \sqrt{-(\tilde{y}^0)^2 + f(\tilde{x}^0)^2((y^1)^2 + (y^2)^2 + (y^3)^2)}, \quad (61)$$

where the conformal factor is nothing but the usual cosmological scale factor and  $\tilde{x}^0$  is the usual cosmological time  $t$ .

For the Finsler function (21) we employ the coordinate change  $\frac{\partial \tilde{x}^0}{\partial x^0} = f(x^0)$ , implying that  $\tilde{y}^0 = f(x^0)y^0$ , so that

$$F = \left( |\tilde{y}^0| + \text{sgn}(\phi) f(\tilde{x}^0) \sqrt{|\phi|} \right) e^{\frac{\tilde{y}^0}{\tilde{y}^0 + \text{sgn}(\phi) f(\tilde{x}^0) \sqrt{|\phi|}}}. \quad (62)$$

Hence, as in the classical FLRW geometry case, the conformal factor can be interpreted as scale factor of the spatial universe,  $x^0$  as conformal time  $\eta$  and  $\tilde{x}^0$  as cosmological coordinate time  $t$ . For now we will adopt this classical cosmology notation.

To be a solution of the Finsler gravity equations we found that  $f(\eta) = c_1 e^{\eta c_2}$ , which implies from the coordinate change between  $\eta$  and  $t$  that

$$dt = c_1 e^{\eta c_2} d\eta \Leftrightarrow \eta(t) = \frac{1}{c_2} \ln \left( \frac{c_2}{c_1} (t - c_3) \right), \quad (63)$$

where  $c_3$  is a constant of integration. Thus, in cosmological time, the scale factor of the vacuum Finsler cosmology we find is

$$f(t) = c_2 (t - c_3). \quad (64)$$

Interestingly, it turns out that these solutions are not only Ricci-flat and conformally flat (by their explicit form), but flat, in the sense that all components of the non-linear curvature tensor  $R^a{}_{bc} = \delta_b N_c^a - \delta_c N_a^b$  vanish. Nevertheless the spacetime has non-trivial geometric features.

## V. DISCUSSION

The solutions that we have presented above are, to the best of our knowledge, the first non-Berwaldian exact solutions to Pfeifer and Wohlfarth's field equation. Known exact solutions are scarce since in particular the Landsberg tensor terms in the field equations are difficult to understand. Employing a unicorn Ansatz, i.e. non-Berwaldian Landsberg spaces, makes our solutions particularly special. We have shown that there is a subclass of our solutions for which the

lightcone structure is physically viable. In fact, it is equivalent to the lightcone of the flat special relativistic Minkowski metric. Moreover, we have shown that one of the solutions has cosmological symmetry, i.e. it is spatially homogeneous and isotropic. Additionally, it is conformally flat, with conformal factor depending only on the timelike coordinate, and we have shown that this conformal factor can be interpreted as the scale factor, which then turns out to be a linear function of cosmological time, leading to the natural interpretation of a linearly expanding (or contracting) Finslerian universe.

As an additional curiosity that we have found that the requirement of a physically light cone structure does not strictly speaking necessitate Lorentzian signature, as is widely assumed. This is illustrated by one of our solutions, which has positive definite signature, and yet has a light cone that is equivalent to the lightcone of flat Minkowski space. It is interesting and surprising that such things are apparently possible in Finsler geometry, and this paper shows the first explicit example of a (positive definite) Finsler metric with this property, which seems to be closely related to lack of smoothness of the Finsler metric in certain nontrivial subsets of  $TM$ .

The results obtained in this paper motivate us to begin a systematic search for cosmological Landsberg spacetimes that solve the field equations, using recent results characterizing cosmological symmetry in Finsler spacetimes [43] and Elgendi's machinery for constructing unicorns using conformal transformations [17, 44]. Since (properly Finslerian) cosmological solutions of Berwald type are necessarily static [43] any interesting such Landsberg spacetime must necessarily be a unicorn.

The exciting next step in the study of Finsler gravity, is to study unicorn solutions of the field equation sourced by the 1-particle distribution function of a kinetic gas, in homogeneous and isotropic symmetry. This scenario describes a realistic universe, filled with a kinetic gas with a non-trivial velocity distribution. As we already obtained a non-trivial solution for vacuum Finsler unicorn cosmology, more realistic, matter sourced solutions will help us to further investigate the conjecture that an accelerated expansion of the universe is caused by the contribution of the velocity distribution of the cosmological gas, which sources a Finslerian spacetime geometry.

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